

Extending the sensitivity limits of VERA/KaVA using inverse phase referencing

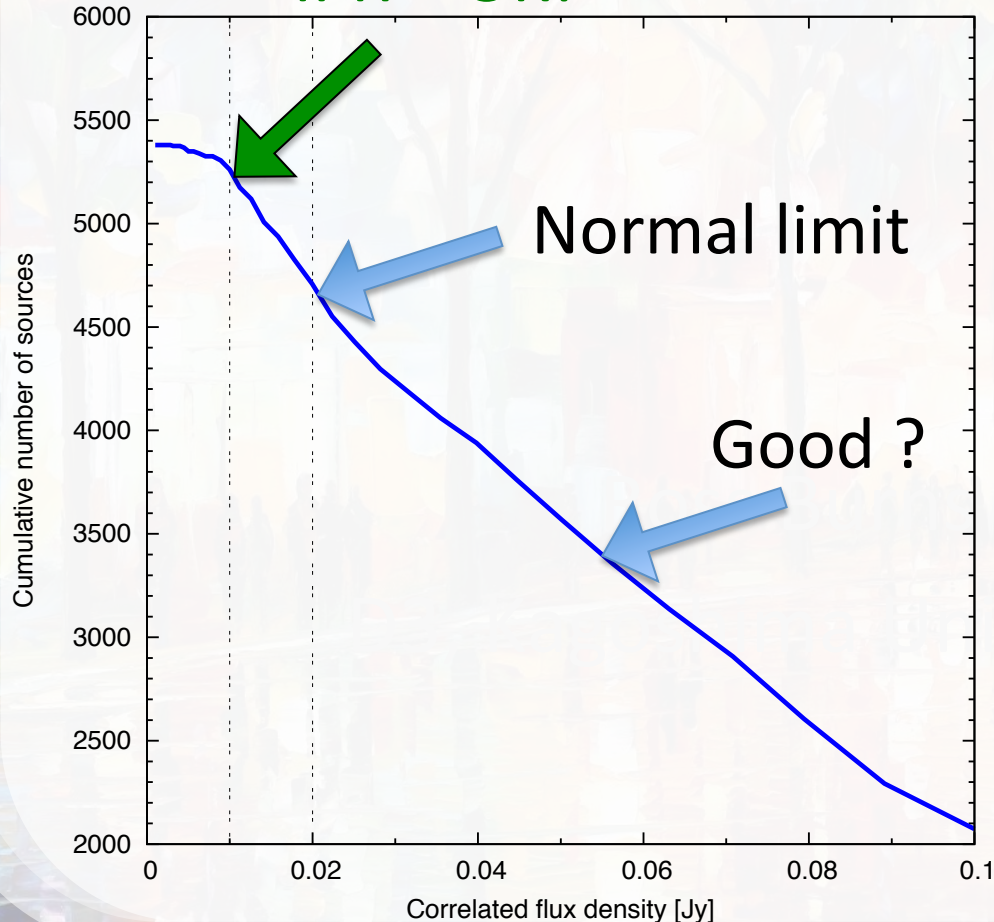
Ross Burns

D3, Kagoshima University

Why?

Known QSOs (X-band)

IPR – OK!



There are more weak QSOs on the sky.

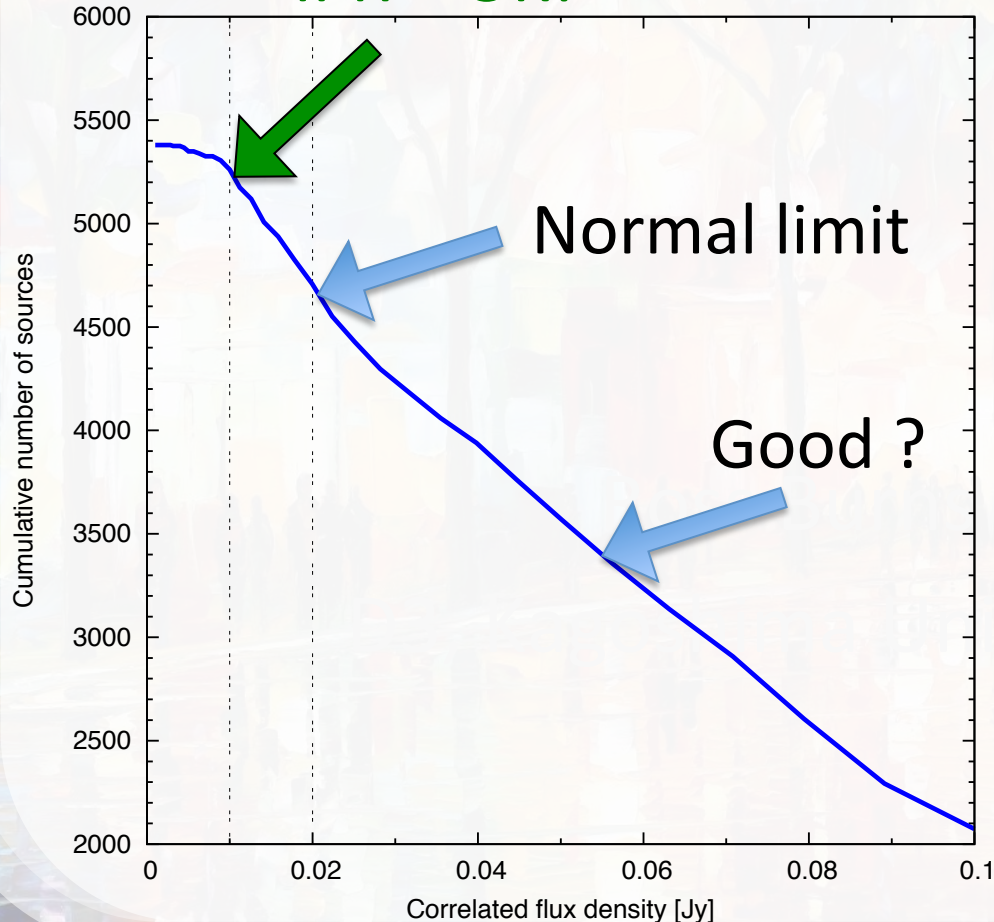
IPR, use weak QSOs

More reference sources available.

Why?

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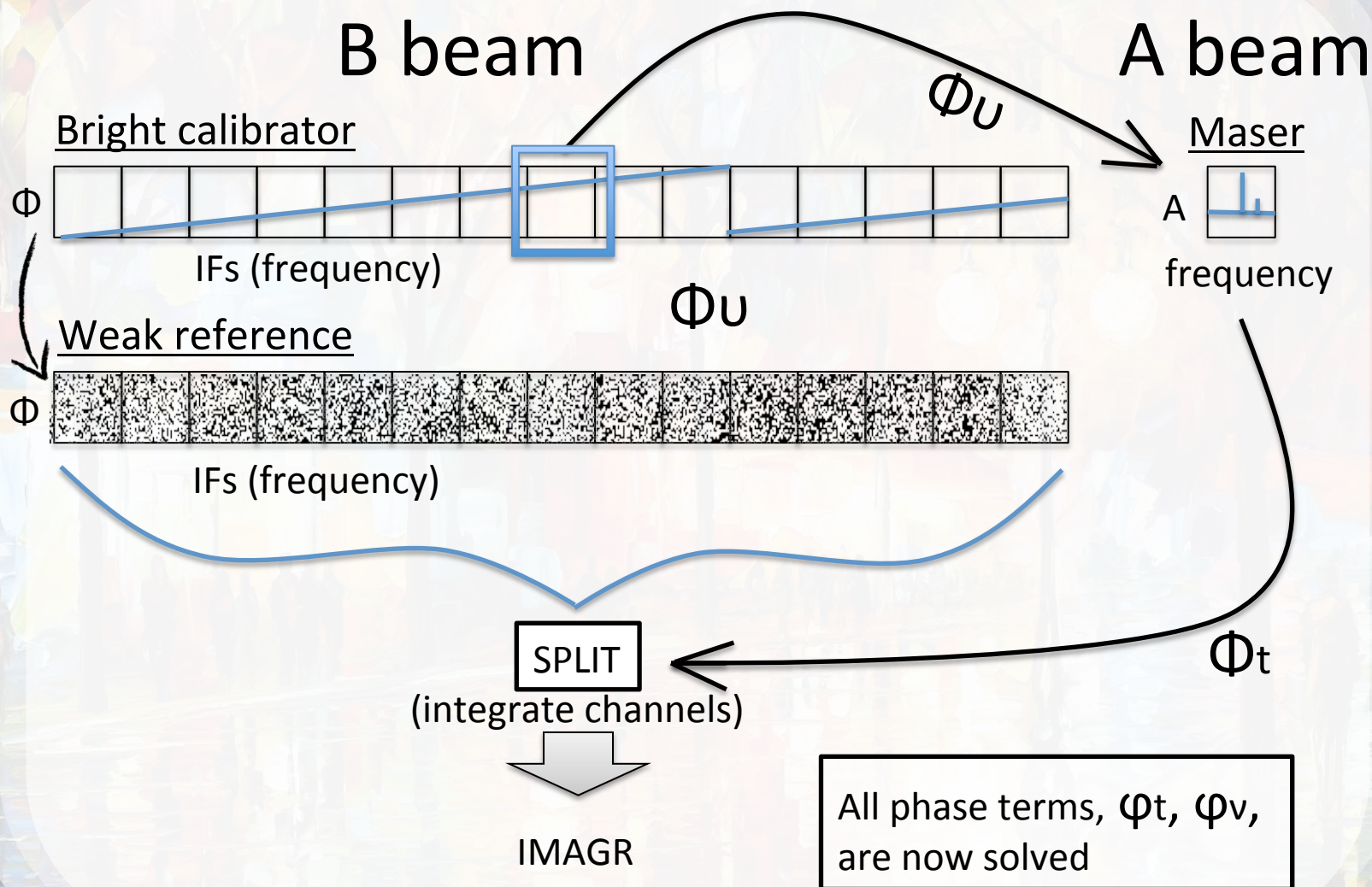


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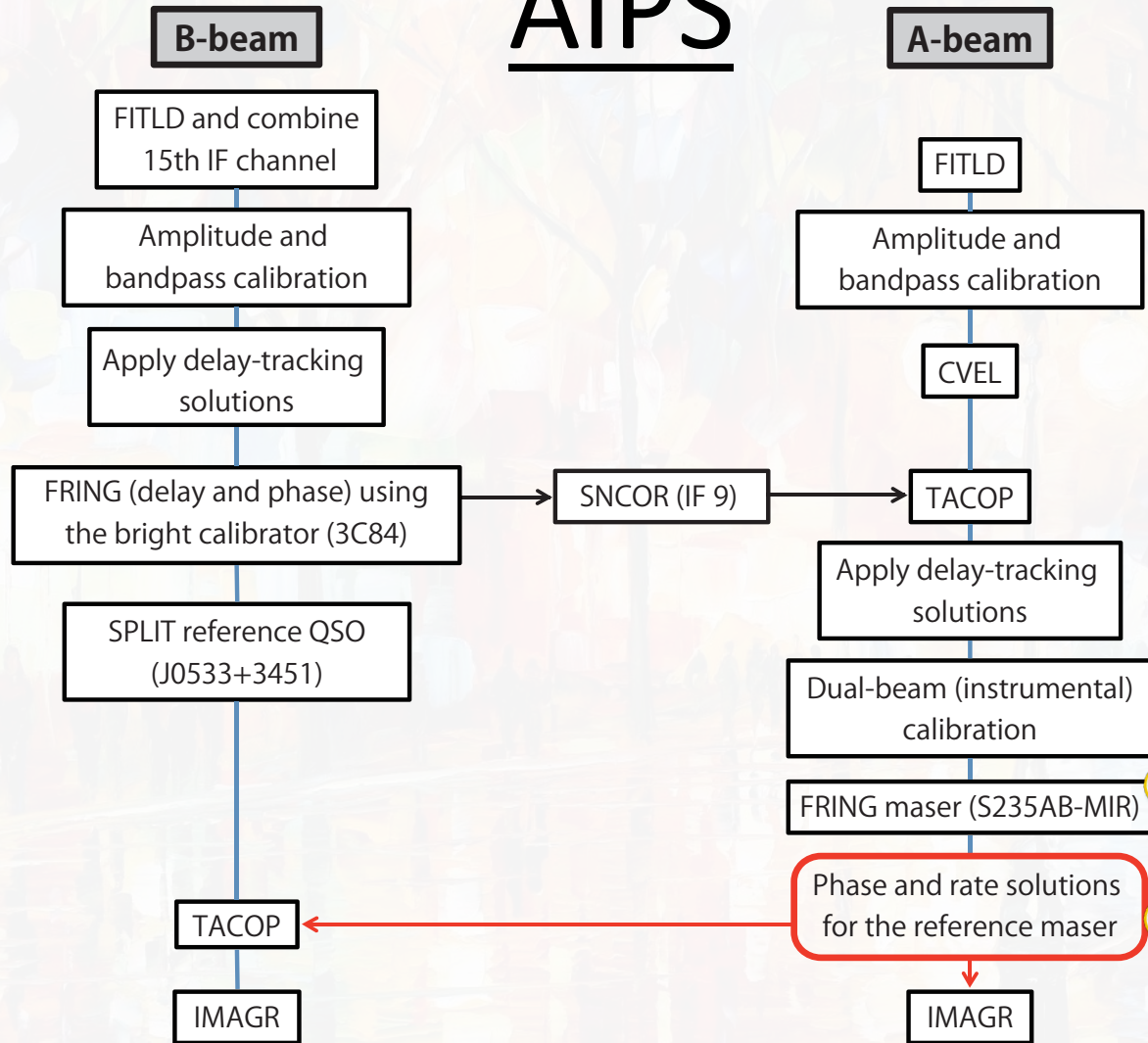
IPR, use weak QSOs

More reference sources available.



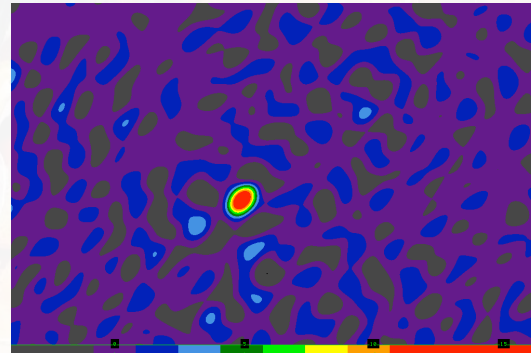
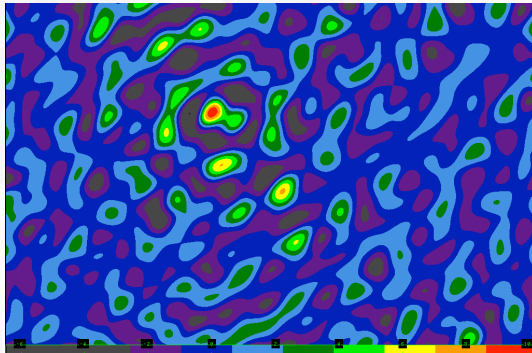


AIPS

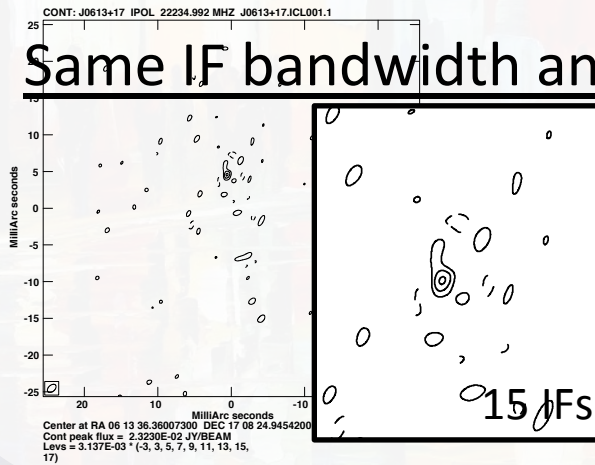


Special points

- Maser must be near the phase tracking center



- Same IF bandwidth and freq. resolution in A and B beams



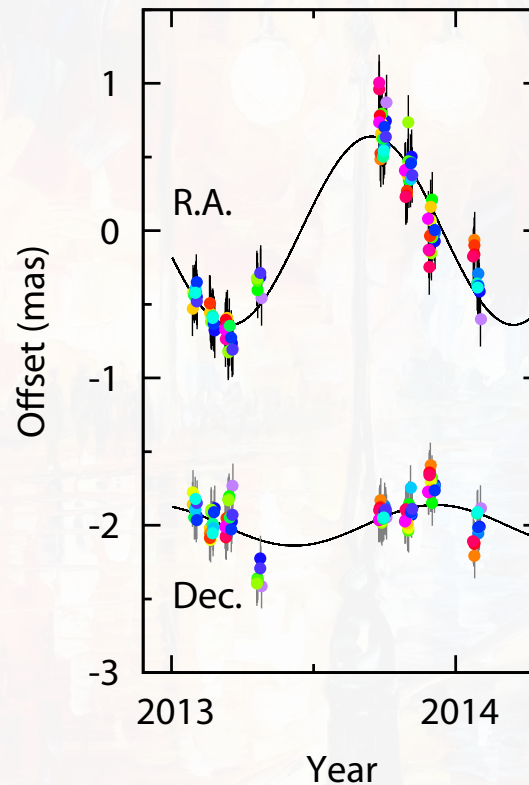
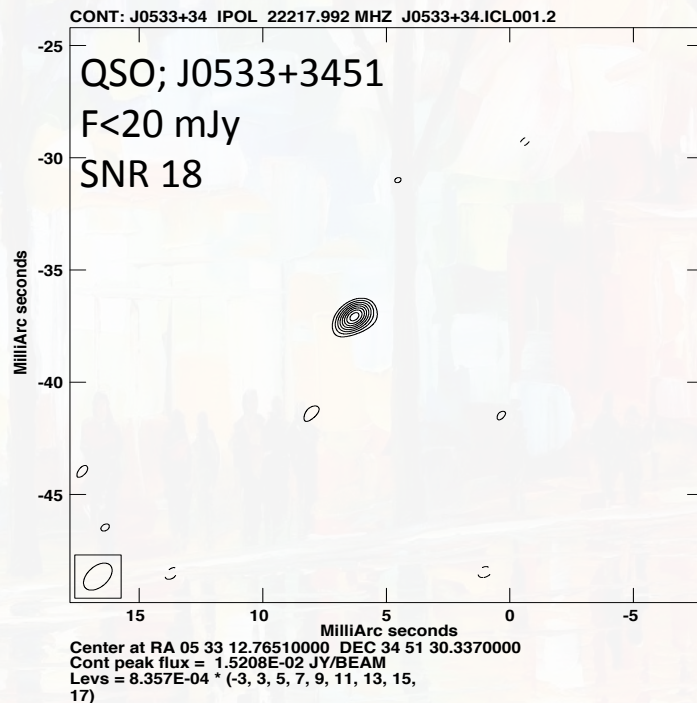
Merits of IPR

- Can use **weak quasars** (detection / non-detection)
More available QSOs
Some masers are ONLY near weak QSOs
Effectively, VERA is more sensitive
- Typically get **more masers** in final images
Maser maps are effectively self-calibrated, without losing absolute astrometry.
(Also, don't lose masers due to coherence loss during phase referencing)
- Faster!! *
Time for FRING is a few seconds

* After re-making "UVW tables"

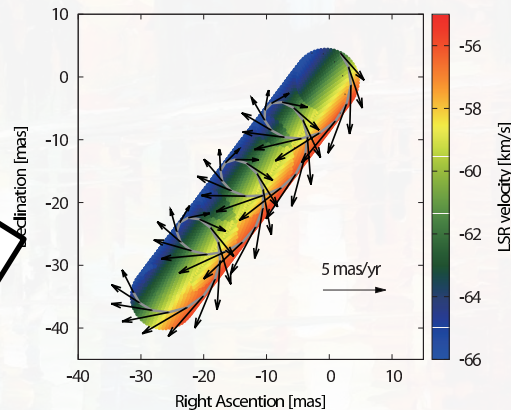
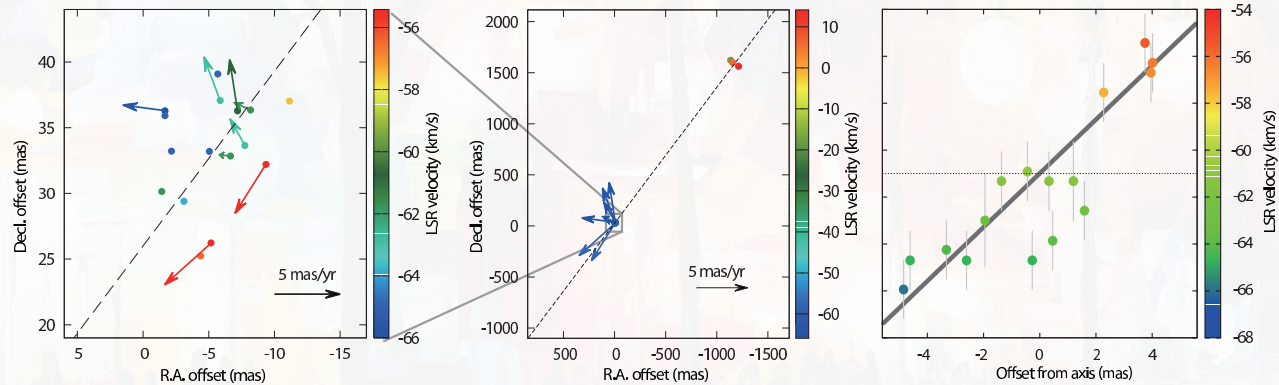
Recent results:

S235AB-MIR



Recent results:

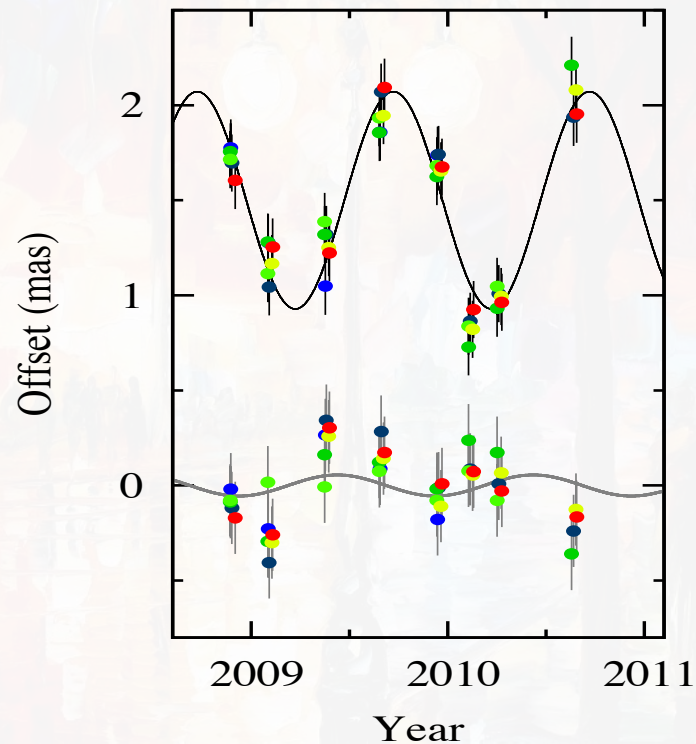
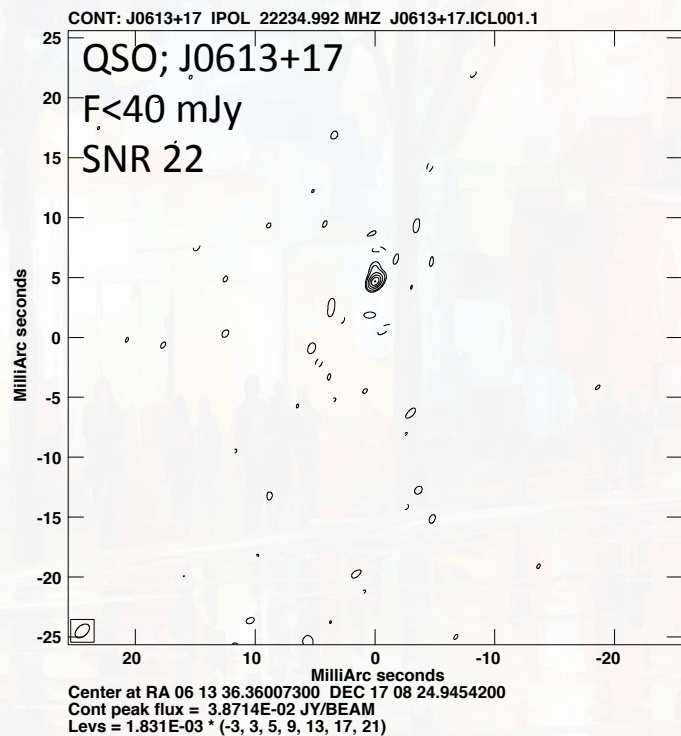
S235AB-MIR



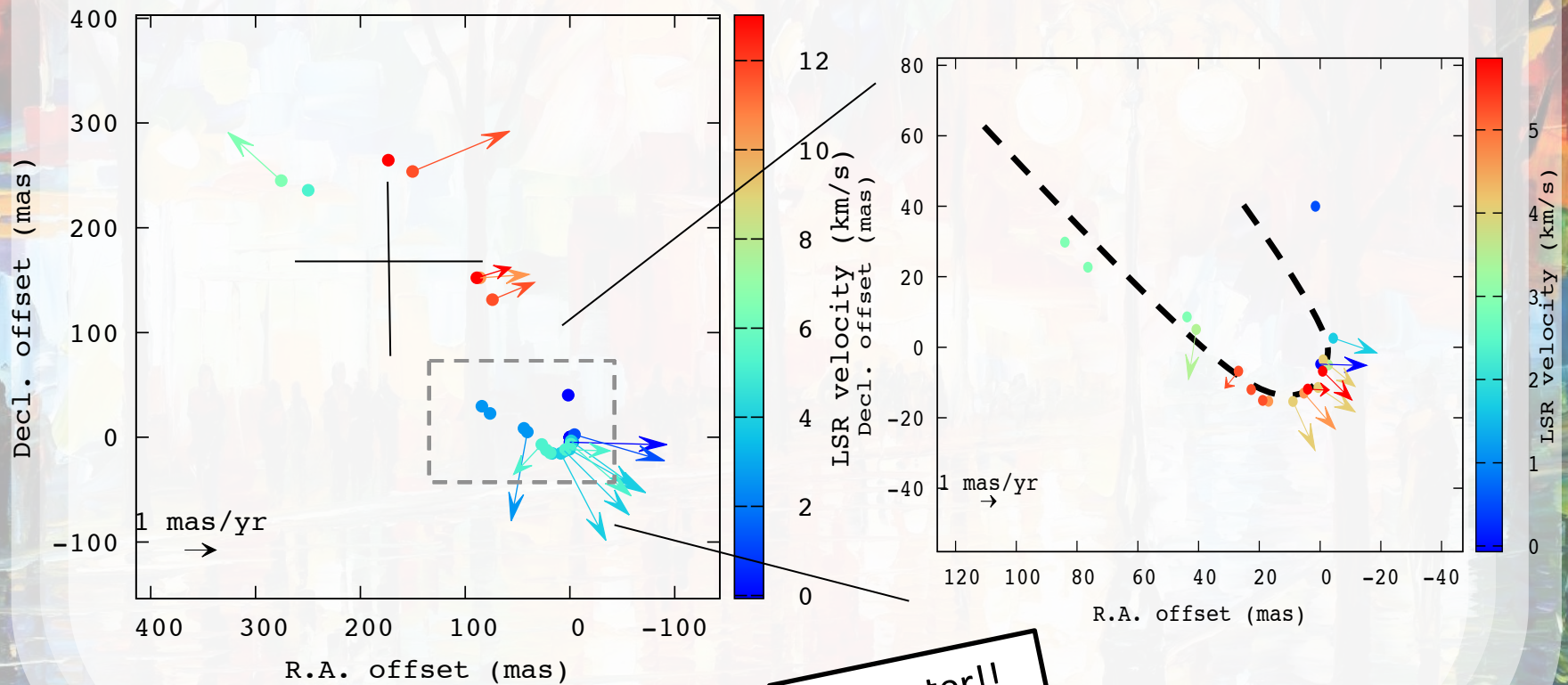
See poster!!



Recent results: S255 (IC2162)



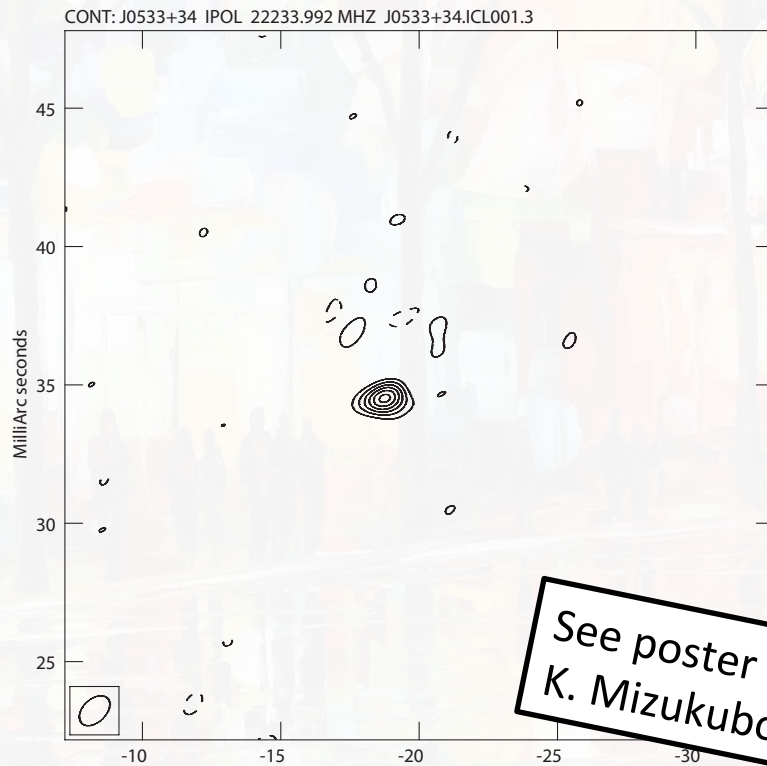
Recent results: S255 (IC2162)



See poster!!

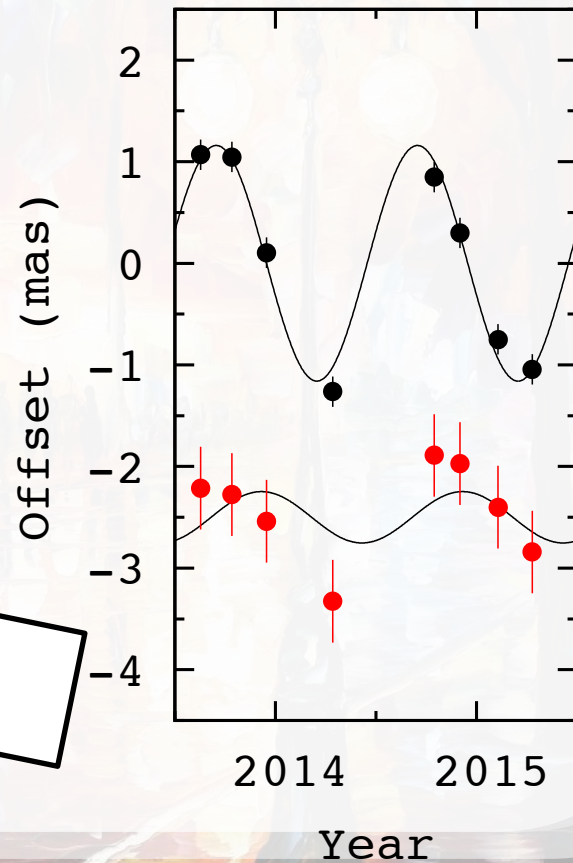
Recent results:

IRAS 05358 – by Mizukubo Kohei



Center at RA 05 33 12.76511100 DEC 34 51 30.3370100
Cont peak flux = 1.4622E-02 JY/BEAM
Levs = 1.062E-03 * (-3, 3, 5, 7, 9, 11, 13, 15)

See poster by
K. Mizukubo !!



Inverse phase referencing AIPS POPS scripts

Full scripts are available on my website:

<http://milkyway.sci.kagoshima-u.ac.jp/~rossburns88/Scripts.html>

Thanks for listening !!